

# Planar infall of CH<sub>3</sub>OH gas around Cepheus A HW2

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## ABSTRACT

**Aims.** In order to test the nature of an (accretion) disk in the vicinity of Cepheus A HW2, we measured the three-dimensional velocity field of the CH<sub>3</sub>OH maser spots, which are projected within 1000 au of the HW2 object, with an accuracy on the order of 0.1 km s<sup>-1</sup>.

**Methods.** We made use of the European VLBI Network (EVN) to image the 6.7 GHz CH<sub>3</sub>OH maser emission toward Cepheus A HW2 with 4.5 milliarcsec resolution (3 au). We observed at three epochs spaced by one year between 2013 and 2015. During the last epoch, in mid-March 2015, we benefited from the newly deployed Sardinia Radio Telescope.

**Results.** We show that the CH<sub>3</sub>OH velocity vectors lie on a preferential plane for the gas motion with only small deviations of 12° ± 9° away from the plane. This plane is oriented at a position angle of 134° east of north, and inclined by 26° with the line of sight, closely matching the orientation of the previously reported disk-like structure. Knowing the orientation of the equatorial plane, we can reconstruct a face-on view of the CH<sub>3</sub>OH gas kinematics onto the plane. CH<sub>3</sub>OH maser emission is detected within a radius of 900 au from HW2, and down to a radius of about 300 au, the latter coincident with the extent of the dust emission at 0.9 mm. The velocity field is dominated by an infall component of about 2 km s<sup>-1</sup> down to a radius of 300 au, where a rotational component of 4 km s<sup>-1</sup> becomes dominant. We discuss the nature of this velocity field and the implications for the enclosed mass.

**Conclusions.** These findings directly support the interpretation that the high-density gas and dust emission that surrounds Cepheus A HW2 traces an accretion disk.

**Key words.** ISM: kinematics and dynamics – masers – stars: formation – stars: individual: Cepheus A HW2

## 1. Introduction

Cepheus A is the second nearest high-mass star-forming region (after Orion), located at a trigonometric distance of 700 pc from the Sun (Moscadelli et al. 2009; Dzib et al. 2011) in the Cepheus OB3 complex (e.g., Fig. 27 of Kun et al. 2008). The region has an IR luminosity in the range  $2\text{--}3 \times 10^4 L_{\odot}$  (e.g., Mueller et al. 2002; De Buizer et al. 2017) and hosts a tight (<7000 au) cluster of young stellar objects (YSOs) associated with radio continuum emission. Half of the bolometric luminosity is attributed to the brightest radio source in the field, named “HW2” after Hughes & Wouterloot (1984, see also Garay et al. 1996), which would correspond to a zero-age main-sequence (ZAMS) star with an early-B spectral type and a mass in excess of 10  $M_{\odot}$ . The HW2 object drives one of the best examples of radio thermal jets in the literature (e.g., Rodriguez et al. 1994; Curiel et al. 2006), which excites several shocked layers of H<sub>2</sub>O maser emission (e.g., Torrelles et al. 2011). At the origin of the radio jet, Patel et al. (2005) resolved a dense core of dust and gas in the 345 GHz band of the Submillimeter Array (SMA; see also Torrelles et al. 2007). The core emission is flattened in the direction perpendicular to the radio jet, with an aspect ratio near 2 between its major and minor axes. Assuming the core is tracing a circumstellar disk, Patel et al. showed that the aspect ratio of the core emission implies a disk inclination of 28° with the line of sight. On the plane of the sky, the major axis of the disk, with an

outer radius larger than 0.8'' (or 560 au), is oriented at a position angle (PA) of 121° (east of north). Although this simple picture would support a scenario of disk-mediated accretion onto HW2, the environment around HW2 is significantly more complex than that of an isolated YSO, showing perhaps more than five objects within a radius of 1000 au (e.g., Fig. 1 of Jiménez-Serra et al. 2009). Because of this multiplicity, the disk scenario has been questioned as due to a chance superposition of different hot cores (e.g., Brogan et al. 2007; Comito et al. 2007).

By directly measuring the velocity field of gas belonging to the putative disk, we can provide a critical test to the disk-like scenario. Compact maser emission centers, on the order of a few au in size, are ideal test particles to probe the local three-dimensional gas kinematics by combining their proper motion vectors (i.e., their displacement on the plane of the sky) with the Doppler shift of the maser lines (e.g., Moscadelli et al. 2011; Torrelles et al. 2011). Toward Cepheus A, bright CH<sub>3</sub>OH maser emission at 6.7 GHz was first reported by Menten (1991) and imaged at milliarcsecond resolution by Sugiyama et al. (2008a,b) and Torstensson et al. (2008). Individual CH<sub>3</sub>OH masers are projected within 1000 au of the HW2 object, and outline a filamentary distribution with local standard of rest (LSR) velocities redshifted by less than 5 km s<sup>-1</sup> with respect to the systemic velocity of the region (around  $-4.5$  km s<sup>-1</sup>). Following this evidence, it has been proposed that the CH<sub>3</sub>OH masers are tracing a contracting circular ring with radius of about 680 au,

**Table 1.** Summary of EVN observations (code ES071).

Epochs <sup>a</sup> (gg/mm/aa)	Run (h)	$\nu_{\text{rest}}$ (MHz)	$\Delta\nu$ (km s <sup>-1</sup> )	$V_{\text{LSR}}$ (km s <sup>-1</sup> )	Absolute position ( $\pm 1$ mas)		HPBW (mas)	rms (mJy beam <sup>-1</sup> )
					RA (J2000) (h m s)	Dec (J2000) ( $^{\circ}$ ' '')		
01/03/13–28/02/14–13/03/15	8	6668.5	0.04	-1.73	22:56:17.9089	62:01:49.527	4.5	4–6

**Notes.** Columns 1 and 2: observing dates and duration of each run. Columns 3 and 4: rest frequency and velocity resolution of the maser lines. Columns 5–7: local standard-of-rest velocity and measured absolute position of the reference maser spot used for calibration. Column 8: restoring beam size (round) used at each epoch. Column 9: thermal noise achieved per line-free channel map at each epoch (Stokes  $I$ ). <sup>(a)</sup> The EVN operated with 8 antennas at the first two epochs (EF, WB, JB, ON, MC, NT, TR, YS), with the addition of SR during the last epoch.

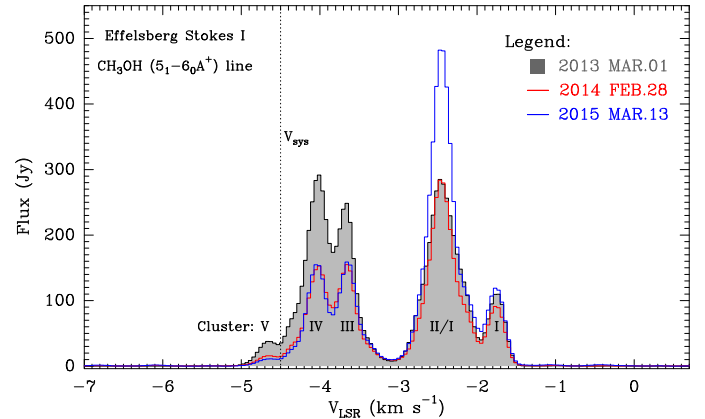
centered on HW2, and oriented nearly edge-on with respect to the observer (Torstensson et al. 2008, 2011; Vlemmings et al. 2010; Sugiyama et al. 2014). Based on the detection of an ordered polarization field through the CH<sub>3</sub>OH maser emission, Vlemmings et al. (2010) also suggested that the maser motions might be driven by magnetic forces.

In order to test the disk-like scenario, we exploited the high sensitivity of the European VLBI Network (EVN) to observe the 6.7 GHz CH<sub>3</sub>OH masers in HW2 at different epochs, with the aim of eventually tracing proper motions with an accuracy of 0.1 km s<sup>-1</sup>. We present the details of the EVN observations in Sect. 2. In Sect. 3 we report the proper motion measurements and the kinematic properties of the maser distribution. In Sect. 4 we model the full-space velocity field of the CH<sub>3</sub>OH gas, and discuss the implications for the disk-like scenario in Cepheus A HW2. Conclusions are drawn in Sect. 5.

## 2. Observations and calibration

We performed multi-epoch phase-referencing EVN observations of the 5<sub>1</sub>–6<sub>0</sub> A<sup>+</sup> CH<sub>3</sub>OH maser transition toward Cepheus A HW2 and correlated all four polarization combinations (RR, LL, RL, and LR). At each epoch, we observed three calibrators together with the maser target. First, the reference polarization calibrator, J1331+3030, was observed at the beginning of the experiment to register the systematic rotation of the linear polarization angle ( $\chi_{\text{pol}}$ ) in the EVN dataset. Second, J2202+4216 was observed as a fringe-finder calibrator every hour to correct both the instrumental phase delay and the polarization leakage. Third, we observed the reference position calibrator J2302+6405, which is in the International Celestial Reference Frame (ICRF) catalog, and is offset by 2° from Cepheus A. Scans on the maser target and the position calibrator were alternated over a cycle of 5 min. The EVN data were processed with the SFXC software correlator (Keimpema et al. 2015) at the Joint Institute for VLBI in Europe by using an averaging time of 2 s and two frequency setups. A high spectral sampling (0.98 kHz) over a narrow band (2 MHz) was used to accurately sample the maser line width. A coarse spectral sampling (15.6 kHz) over a wide band of 16 MHz was used to improve on the continuum sensitivity of the calibrators maps. The observation information is summarized in Table 1.

Data were reduced with the NRAO Astronomical Image Processing System (AIPS) following standard procedures. Fringe fitting and self-calibration were performed on a strong maser channel at an LSR velocity of -1.73 km s<sup>-1</sup> at each epoch. We produced total intensity (Stokes  $I$ ) maps of the maser emission to image an area of radius 1''5 around HW2. At each epoch, we set a restoring beam size of 4.5 mas (round), equal to the geometrical average of the *clean* beam size obtained with ROBUST 0 weighting (task IMAGR of AIPS). Figure 1 shows the Effelsberg



**Fig. 1.** Effelsberg total-power spectra toward Cepheus A HW2 obtained from the EVN observations at C band. Observing dates are indicated in the top right corner. The dotted vertical line indicates the rest velocity ( $V_{\text{sys}}$ ) of the HW2 object. According to Fig. 2, maser emission associated with the different clusters is also indicated.

spectra of the CH<sub>3</sub>OH maser emission toward Cepheus A HW2 at each epoch.

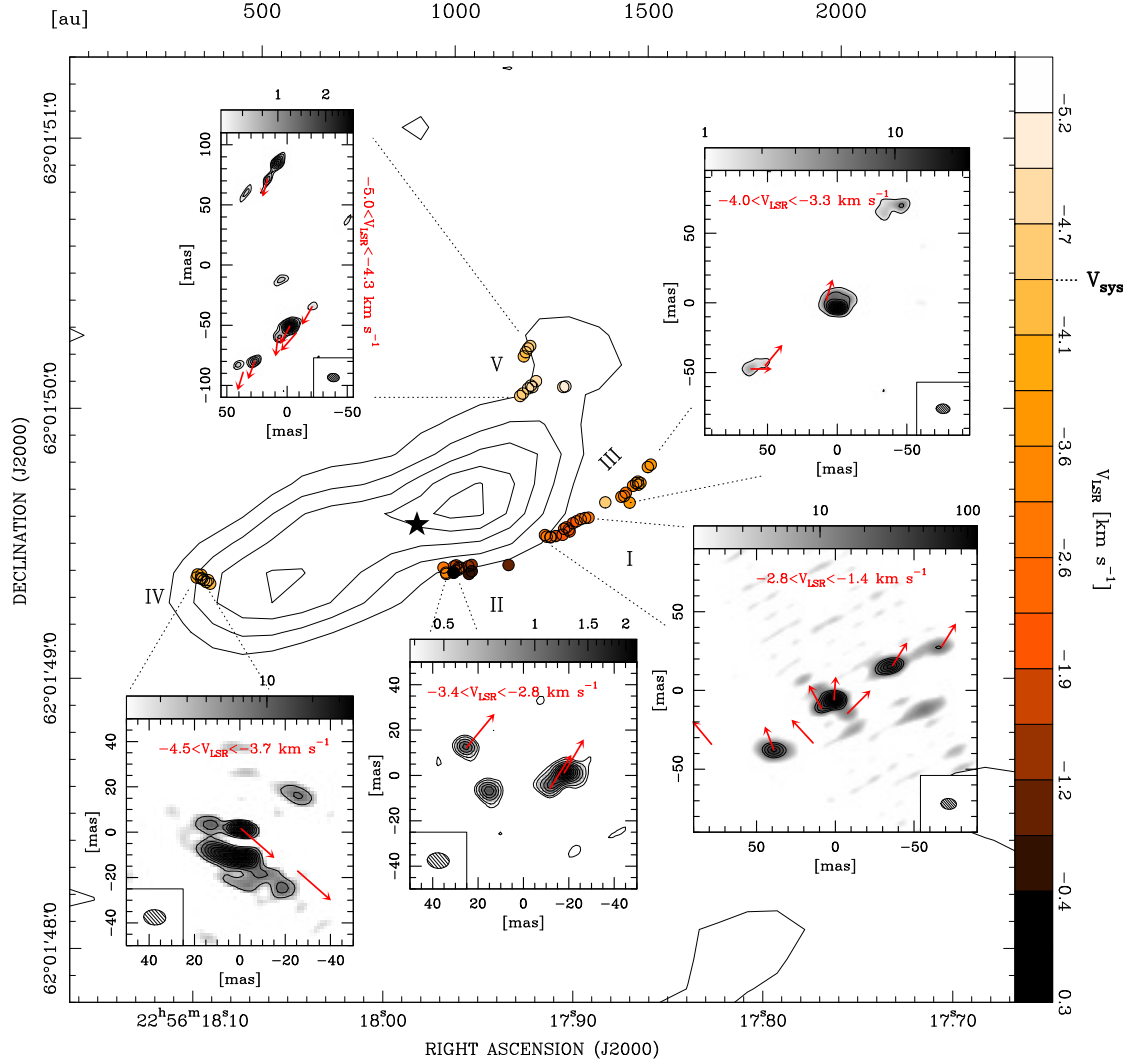
We postpone the analysis of the polarization information to a subsequent paper. Recently, Lankhaar et al. (2016) presented a theoretical derivation of the hyperfine structure of the methanol molecule. They showed that a number of hyperfine transitions, close in frequency, might contribute to the 6.7 GHz CH<sub>3</sub>OH maser emission. In order to account for this multiplicity, a new pumping model for the 6.7 GHz CH<sub>3</sub>OH masers is required. These findings might have an influence on the modeling of the polarized maser emission (Vlemmings et al. 2010), which is under review (Lankhaar et al., in prep.; Vlemmings et al., in prep.).

## 3. Kinematic properties of the maser distribution

We detected 64 individual maser emission centers, or *cloudlets*<sup>1</sup>, above a threshold of  $7\sigma$ . For strong maser channels, this threshold is typically set by the limited dynamical range of the

<sup>1</sup> Hereafter, we make use of the following convention:

- a maser *spot* is a compact emission, at milliarcsecond scale, detected on a single channel map and best-fitted with an elliptical Gaussian brightness distribution;
- a maser *cloudlet*, elsewhere referred to as “feature”, is an individual gas condensation that is composed of a cluster of contiguous spots both in space ( $< \text{HPBW}/10$ , typically) and LSR velocity ( $\leq 1/5 \times$  line width, based on an average line width of 0.2–0.3 km s<sup>-1</sup>).



**Fig. 2.** Distribution of the 6.7 GHz CH<sub>3</sub>OH maser cloudlets detected within  $\sim 1$  arcsec of Cepheus A HW2. The *main panel* shows the distribution of the maser cloudlets (colored dots) superimposed on a HC<sub>3</sub>N (27–26)  $v_7 = 1_e$  map (contours) integrated in the velocity range from  $-7$  to  $-3$  km s $^{-1}$  (from Jiménez-Serra et al. 2009). HC<sub>3</sub>N contours start at  $2\sigma$  by  $2\sigma$  steps. The upper axis gives the linear scale of the map at a distance of 700 pc. Maser colors indicate the gas LSR velocity locally, according to the right-hand wedge. Maser groups are labeled according to Sugiyama et al. (2008b). The systemic velocity ( $V_{\text{sys}}$ ) is set at  $-4.5$  km s $^{-1}$ . The black star marks the HW2 position according to Curiel et al. (2006). *Insets*: blow-up of the local maser emission (beam size at a bottom corner) summed within the  $V_{\text{LSR}}$  range specified in each box (see Sect. 3). The gray logarithmic scale (top wedge) gives the brightness scale from  $10\sigma$  up to the peak emission in Jy beam $^{-1}$ ; the black contours give the brightness levels from 10% of the peak emission by 10% steps. The red arrows trace the proper motion vectors of the maser cloudlets; the arrow length indicates the position reached after a time of 40 yrs at the average velocity of  $1.7$  km s $^{-1}$ . Note that a few proper motions are associated with faint maser cloudlets that are not visible in the summed maps.

images. In Table A.1 we list the relative position of each maser cloudlet with respect to the peak of the reference maser channel (used to calibrate the visibilities). The absolute position of this maser spot is determined with an accuracy of  $\pm 1$  mas and is given in Table 1. The centroid position of each CH<sub>3</sub>OH cloudlet is determined by an intensity-weighted average of the spots' distribution within a beam, following Sanna et al. (2010). For cloudlets brighter than  $1$  Jy beam $^{-1}$ , positional uncertainties are typically better than  $0.1$  mas.

Maser cloudlets are spread within a projected radius of  $1000$  au from the HW2 object and are arranged in five clusters by position (Fig. 2). These clusters have been labeled, from I to V, after Sugiyama et al. (2008a). Clusters I and II, which project closer to the HW2 object, show the most redshifted emission and span a range of LSR velocities between  $0$  and  $-3.5$  km s $^{-1}$ . The three clusters III, IV, and V, projected farther away from HW2, are centered on LSR velocities (intensity-weighted) of  $-3.63$ ,

$-4.05$ , and  $-4.54$  km s $^{-1}$ , respectively, and each cluster emits over a narrow range of velocities (ca.  $0.7$  km s $^{-1}$ ). The five insets of Fig. 2 show the details of the maser emission, summed over the relevant LSR velocity ranges, at the position of each cluster (from the first epoch data). These maps were produced after a boxcar smoothing of three velocity channels of the initial dataset, and by imaging the new dataset with an  $u-v$  tapering of  $15$  M $\lambda$ . At variance with the integrated maser flux, which shows significant variations among the epochs (up to a factor of 2), the relative positions among the clusters and their ranges of LSR velocities have remained constant during the past 10 yrs, since the first high-resolution observations by Torstensson et al. (2011, their Fig. 1) in November 2004.

In order to provide proper motion measurements accurate to a few  $0.1$  km s $^{-1}$ , we studied the spatial and spectral distribution of the maser spots within each cloudlet detected at the three epochs (e.g., Sanna et al. 2010, their Fig. 6). At a distance

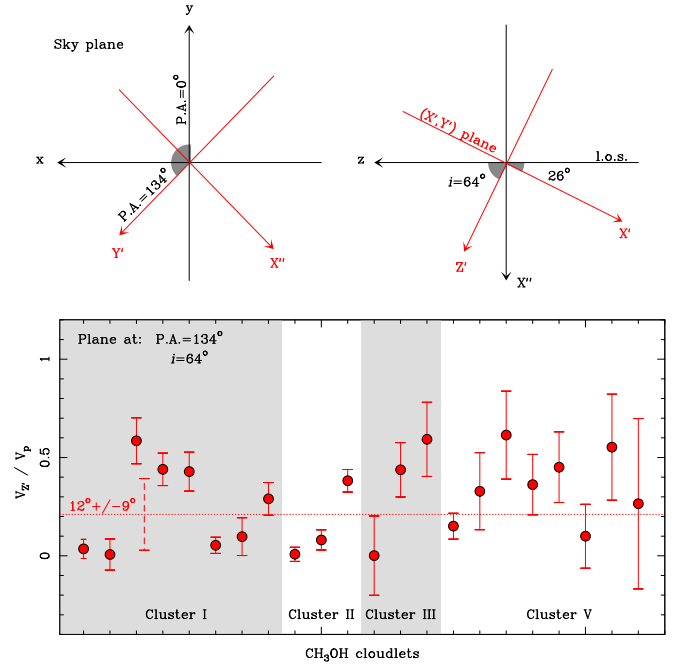
of 700 pc, a proper motion of  $1 \text{ km s}^{-1}$  corresponds to a displacement of 0.6 mas every two years. We do not report proper motions for cloudlets that changed their internal spots' distribution over the epochs so that their centroid position shifted by more than 0.3 mas. Given that we expect proper motions of a few  $\text{km s}^{-1}$ , based on the LSR velocity dispersion of the maser emission, this criterion allows us to exclude proper motions that are severely affected by the internal cloudlet structure. We selected a subsample of 24 maser cloudlets that satisfy this condition and calculated their relative proper motion vectors with respect to the reference cloudlet at zero offset (number 2) by interpolating their positions among the epochs with a linear regression fit (e.g., [Moscadelli et al. 2010](#), their Fig. 3). Proper motions were referred to the reference system of the gas at rest by subtracting the average velocity vector of the whole sample. In Cols. 6 and 7 of Table A.1 we list the eastward and northward velocity components of the corrected proper motions. Their magnitude does not exceed  $5 \text{ km s}^{-1}$  and has an average value of  $1.7 \text{ km s}^{-1}$ . The mean uncertainty per velocity component is  $0.2 \text{ km s}^{-1}$ . In each inset of Fig. 2, we plot the direction of the proper motion vectors at each measured position (red arrows). The proper motion magnitude gives the position that would be reached by a maser cloudlet, moving at an average velocity of  $1.7 \text{ km s}^{-1}$  after a time of 40 yr.

## 4. Discussion

### 4.1. Preferential plane of gas motion

Because  $\text{CH}_3\text{OH}$  molecules are (mainly) formed on grain surfaces and are evaporated in the gas phase at a temperature (100–300 K; [Herbst & van Dishoeck 2009](#)) that is well below the dust sublimation temperature ( $\sim 1500 \text{ K}$ ), we expect that the presence of  $\text{CH}_3\text{OH}$  gas should be coupled with the presence of dust condensation. The spatial distribution of dust continuum emission and high-density molecular lines toward Cepheus A HW2 has been interpreted as the silhouette of an accretion disk with a radius of several 100 au. As a comparison, in Fig. 2 we overlap the distribution of the  $\text{CH}_3\text{OH}$  maser cloudlets with an integrated velocity map of the  $\text{HC}_3\text{N}$  emission in the  $v_7 = 1$  excited state (from [Jiménez-Serra et al. 2009](#)). The  $\text{CH}_3\text{OH}$  maser transition at 6.7 GHz is inverted by IR radiation in the range 20–30  $\mu\text{m}$ , which is enhanced in the dusty environment heated by a young massive star (e.g., [Cragg et al. 2005](#)). Similarly, the  $\text{HC}_3\text{N}$  vibrationally excited states are populated by radiative pumping that is due to an IR field at wavelengths of a few 10  $\mu\text{m}$  ([Goldsmith et al. 1982](#)). Figure 2 emphasizes the spatial correlation between the dense molecular envelope surrounding HW2 and the loci of  $\text{CH}_3\text{OH}$  maser emission.

Having on hand the information of the full-space motion of the  $\text{CH}_3\text{OH}$  gas, we searched for a preferential plane of gas motion as would be expected in the presence of a disk. With reference to the upper panel of Fig. 3, we made two consecutive rotations of the original reference system in equatorial coordinates ( $\text{RA} \equiv x, \text{Dec} \equiv y$ ): 1) we projected the maser proper motions on the  $(X', Y')$  reference system, rotated by an angle “PA” on the plane of the sky; 2) taking the LSR velocity components into account as well, we followed a second projection onto the reference system  $(X', Y', Z')$ , obtained by rotating the coordinates  $(X'', Z)$  around the axis  $Y'$  by an angle “ $i$ ”. We iteratively explored the full range of  $(X', Y')$  plane orientations varying PA and  $i$  in the ranges 0–360° and 0–90°, respectively. At each run, we calculated the ratio between the maser velocity components

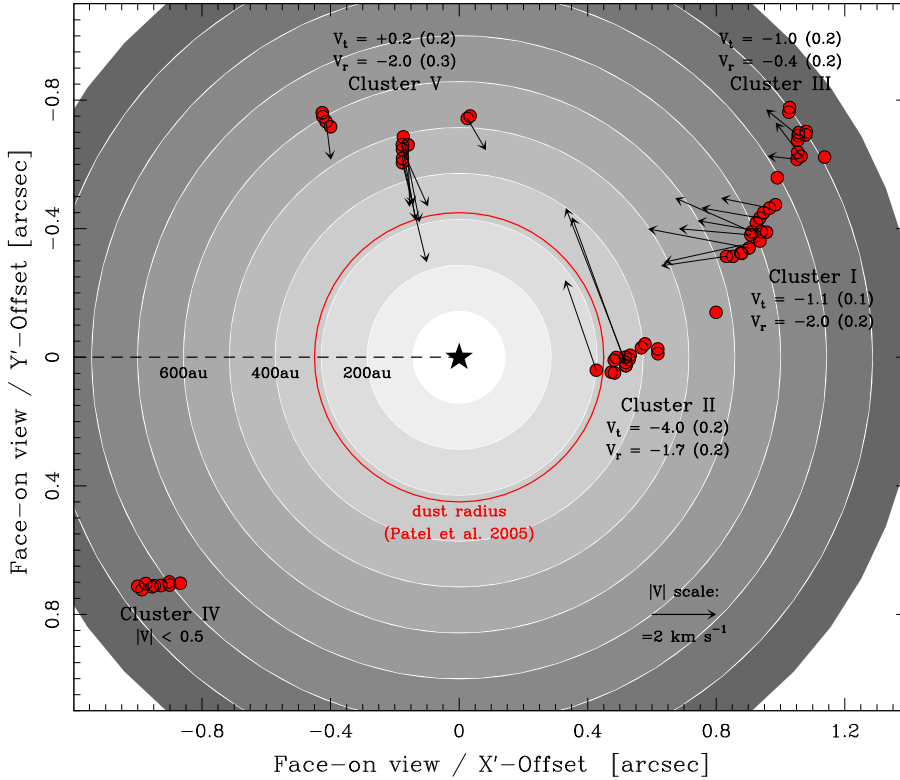


**Fig. 3.** Orientation of the best plane  $(X', Y')$  that minimizes the velocity component ( $V_{Z'}$ ) of the  $\text{CH}_3\text{OH}$  masers perpendicular to the plane. *Upper panel:* left and right plots show a view of the best plane as projected on the plane of the sky and with respect to the line of sight, respectively. *Lower panel:* ratio of the velocity components perpendicular ( $V_{Z'}$ ) and parallel ( $V_p$ ) to the best plane, plotted for each maser cloudlet (grouped by clusters). Error bars give the  $1\sigma$  uncertainty of the ratio  $V_{Z'}/V_p$ . The dotted line marks the weighted average of  $V_{Z'}/V_p$  (with its uncertainty), which corresponds to an average deviation of  $12^\circ$  away from the best plane.

perpendicular ( $V_{Z'}$ ) and parallel ( $V_p$ ) to the plane  $(X', Y')$ , in order to estimate how well the velocity vectors accommodate the  $(X', Y')$  plane. The upper panel of Fig. 3 shows the orientation of the best plane (in red), which minimizes the weighted average ratio  $V_{Z'}/V_p$  for the maser sample of 24 cloudlets. This plane intersects the plane of the sky along a line at a PA of 134° and it is inclined by 26° with respect to the line of sight. In the lower panel of Fig. 3, we plot the ratio  $V_{Z'}/V_p \pm 1\sigma$  for each maser cloudlet separately (dots and error bars). We do not plot the two masers belonging to cluster IV because they have motions slower than  $0.5 \text{ km s}^{-1}$  and high relative uncertainties. The weighted average of the sample is drawn with a dashed horizontal line in the lower panel of Fig. 3 and corresponds to a residual inclination of  $12^\circ \pm 9^\circ$  with respect to the plane  $(X', Y')$ . Alternatively, this inclination corresponds to an average velocity component of  $0.5 \text{ km s}^{-1}$  away from the plane.

The best-fit plane (hereafter, the equatorial plane) over which  $\text{CH}_3\text{OH}$  masers move is coincident with the orientation of the flattened core observed by [Patel et al. \(2005\)](#) within about  $10^\circ$  (PA and  $i$  of  $121^\circ$  and  $62^\circ$ , respectively). It is also worth noting that the orientation of the radio jet, at a PA near  $45^\circ$  ([Curiel et al. 2006](#), their Table 2), coincides with the minor axis of the equatorial plane as projected onto the sky ( $X''$  in Fig. 3). These findings support the interpretation that the core emission imaged by [Patel et al.](#) belongs to a continuous structure that extends along the equatorial plane of gas motion. In the next section, we analyze the gas motion in the equatorial plane and show that maser cloudlets participate in a global infall toward HW2.





**Fig. 4.** Face-on view of the plane at a PA of 134° and inclination of 64°, which contains the full-space motions of the 6.7 GHz CH<sub>3</sub>OH maser cloudlets (see Fig. 3). The star at the center of the plot pinpoints the position of the HW2 object that powers the radio continuum jet perpendicular to the plane (Curiel et al. 2006). Gray circles mark different distances from the YSO position along this plane. Red dots mark the different locations of the CH<sub>3</sub>OH maser cloudlets, and the black arrows trace their motions in the plane. Each maser cluster is labeled according to Fig. 2. The velocity scale is reported at the bottom right corner. For each cluster, we report the radial ( $V_r$ ) and tangential ( $V_t$ ) group velocities with respect to the YSO position ( $1\sigma$  uncertainty in brackets). The red circle, at a radius of about 300 au, marks the deconvolved size of the dust emission at 0.9 mm observed by Patel et al. (2005).

#### 4.2. Analysis of the planar motion around HW2

Knowing the orientation of the equatorial plane, we can produce a face-on view of the CH<sub>3</sub>OH maser velocities in this plane. In Fig. 4 we plot the position of the CH<sub>3</sub>OH cloudlets (red dots) projected onto the equatorial plane. We set the origin of the plot at the position of HW2 as defined by Curiel et al. (2006). Each cluster of maser emission is labeled according to Fig. 2 for comparison. The result shown in Fig. 4 is independent of any assumption about the maser geometry. Individual centers of maser emission do not originate at a same distance from the star. Maser emission is excited over a range of radii between 300 au and 900 au from HW2. The inner radius coincides with the extent of the dust emission (hereafter, the dust radius) at a wavelength of 0.9 mm (red circle in Fig. 4). A disk radius of about 300 au also marks the boundary where hydrogen density might exceed a threshold of  $10^8 \text{ cm}^{-3}$  (Kuiper et al. 2015, their Fig. 9), above which the CH<sub>3</sub>OH maser transition at 6.7 GHz is quenched (Cragg et al. 2005).

The black arrows in Fig. 4 give the magnitude and direction of the velocity vectors in the equatorial plane. Since maser cloudlets belonging to a given cluster show ordered velocity vectors, we derived a group motion in the radial ( $V_r$ ) and tangential ( $V_t$ ) directions with respect to HW2 for each cluster. These values are reported together with the cluster labels in Fig. 4. A negative radial velocity corresponds to an inward motion toward HW2; a negative tangential velocity indicates a counter-clockwise rotation around HW2. For radii between 300 au and 800 au (clusters I, II, and V), maser velocities have an average inward component of about  $-2 \text{ km s}^{-1}$  within an uncertainty of  $0.2 \text{ km s}^{-1}$ . We explicitly note that the existence of a planar infall toward HW2 in turn supports the existence of an accretion disk. At the loci sampled by the maser emission, the magnitude of the infalling motions appears to be almost constant with the distance from the star within the uncertainties. Outside a radius of 800 au (clusters III and IV), the radial component is more than

four times slower ( $<0.5 \text{ km s}^{-1}$ ) than in the inner regions. On the other hand, there is no clear pattern for the tangential velocity outside a radius of 300 au from the star: clusters I and III have a consistent tangential component of about  $-1 \text{ km s}^{-1}$ , clusters IV and V have a tangential component near zero, and cluster II has a faster tangential component of  $-4 \text{ km s}^{-1}$ . We do not have enough maser points around the star to study the dependence (if any) of the tangential velocity on the azimuthal angle.

#### 4.3. Estimate of the central mass of HW2

Interestingly, the maser cluster (II) closer to the star exhibits a higher value of the tangential velocity. If we assume that this tangential component represents an ordered rotational motion around HW2, we can evaluate the enclosed mass within the dust radius. Assuming Keplerian rotation, for a rotational velocity of  $4 \text{ km s}^{-1}$  and a radius of 315 au, derived from the deconvolved size of the dust continuum map of Patel et al. (2005), we infer a central mass of  $5.5 M_\odot$ . A ZAMS star with a mass of  $6 M_\odot$  and a radius of  $3 R_\odot$  has a bolometric luminosity of  $1 \times 10^3 L_\odot$  (e.g., Schaller et al. 1992), one-tenth of the bolometric luminosity associated with the HW2 object. An additional contribution to the total luminosity might come from accretion. We make use of the infalling motion measured with the masers to estimate the mass infall rate onto the star. In spherical symmetry, for an hydrogen density of  $10^8 \text{ cm}^{-3}$  at a distance of 300 au from the central object, an infall velocity of  $2 \text{ km s}^{-1}$  implies an accretion rate of  $3 \times 10^{-4} M_\odot \text{ yr}^{-1}$ . If the same infall rate reaches the star surface, the accretion luminosity of  $9 \times 10^3 L_\odot$  would dominate the total luminosity. Taking into account that the infalling mass is channeled through the disk onto the star within a small solid angle, the accretion rate might be reduced by a factor  $\sim 0.1$  for a flared disk with an opening angle of  $60^\circ$ . Under these conditions, the total budget of stellar and accretion luminosities would account for a bolometric luminosity of about  $2 \times 10^3 L_\odot$ .

Different lines of evidence suggest an ordered magnetic field in the inner 1000 au of HW2, which might reconcile the kinematic estimate of the stellar mass with a bolometric luminosity on the order of  $1 \times 10^4 L_\odot$ . Recently, [Fernández-López et al. \(2016\)](#) reported CARMA dust polarization observations at 1.3 mm in the vicinity of HW2. The dust polarization vectors align in a direction ( $57^\circ \pm 6^\circ$ ) roughly parallel to the minor axis of the accretion disk on the plane of the sky ( $X''$  in Fig. 3). If the main cause of polarized emission is grain alignment via magnetic forces, [Fernández-López et al.](#) propose a scenario where an uniform magnetic field is threading the accretion disk of HW2. At the scale of individual pockets of gas of a few au in size, the linearly polarized emission of the OH, H<sub>2</sub>O, and CH<sub>3</sub>OH masers still shows coherent polarization vectors either parallel or perpendicular to the same  $X''$  axis of Fig. 3 ([Bartkiewicz et al. 2005](#); [Vlemmings et al. 2006, 2010](#)). The presence of a magnetic field might result in rotational velocities lower than half the expected Keplerian velocity at a given radius, as depicted in Fig. 7 of [Seifried et al. \(2011\)](#). In the case of HW2, if the tangential velocity measured with the masers were (magnetically) reduced by a factor as small as 1/4 of the pure Keplerian rotation, the revised mass estimate would amount to  $10 M_\odot$  and would imply a ZAMS luminosity of about  $1 \times 10^4 L_\odot$  (e.g., Eq. (13) of [Hosokawa et al. 2010](#)).

#### 4.4. Comparison with previous models

In order to explain the CH<sub>3</sub>OH maser distribution around HW2, previous models assumed that the maser emission arises along a circular ring ([Torstensson et al. 2008, 2011](#); [Vlemmings et al. 2010](#); [Sugiyama et al. 2014](#)). Maser positions were fitted by an elliptical ring on the plane of the sky, with free parameters the major ( $2 \times a$ ) and minor axes of the ellipse, the position angle of the major axis ( $PA_{\text{ell}}$ ), and the center of the ellipse. If the ellipse is the projection of a circular ring tilted with respect to the line of sight, the inclination ( $i$ ) of the circular ring is determined from the ratio of the minor-to-major axes of the ellipse. Since the molecular outflow entrained by the radio jet has the blueshifted lobe to the NE and the free-free optical depth of the radio jet favors the brightest masers to be on the near side of the ring, [Torstensson et al.](#) first argued that maser clusters I and II have to lie on the near side of the ring. Under these conditions, the plane of the circular ring would be oriented at a position angle ( $PA_{\text{ell}}$ ) between  $99^\circ$  and  $110^\circ$ , with an inclination ( $i$ ) between  $68^\circ$  and  $73^\circ$ ; the ring radius ( $a$ ) would range between 650 au and 680 au ([Vlemmings et al. 2010](#); [Torstensson et al. 2011](#); [Sugiyama et al. 2014](#)). The position angle of the equatorial plane derived in Sect. 4.1 differs by about  $30^\circ$  with respect to the plane of the circular ring; the radius of the circular ring approximates the average distance of the maser cloudlets with respect to HW2 (Fig. 4).

To explain the three-dimensional magnetic field configuration around HW2 as inferred from the 6.7 GHz maser emission, [Vlemmings et al.](#) also proposed that the maser cloudlets might be detached by a few 100 au from the equatorial plane. Maser cloudlets would move along the magnetic field lines at the edge of the disk, where the lines warp toward the central core (Fig. 6 of [Vlemmings et al. 2010](#)). This picture might still hold even if the maser emission does not arise at a constant radius, since the toroidal magnetic field lines would thread the disk plane at different radii from the central star. The current analysis does not rule out that the maser cloudlets may lie at small heights ( $h$ ) above the equatorial plane ( $h \lesssim 100$  au), provided that the velocity field is still mainly planar (e.g., Fig. 9 of [Seifried et al. 2011](#)).

[Torstensson et al.](#) compared the line-of-sight velocities of the maser emission with those expected for a circular ring either in Keplerian rotation or with a dominant radial motion. They inferred negligible Keplerian rotation and proposed that the ring itself is contracting toward HW2 with a velocity on the order of  $1 \text{ km s}^{-1}$ . [Sugiyama et al.](#) also measured the 6.7 GHz maser proper motions with the Japanese VLBI Network (JVN). They modeled the three-dimensional velocity field of the CH<sub>3</sub>OH gas as the composition of a radial and tangential velocity with respect to the ring, and derived a global infalling motion of  $2.0 \pm 0.6 \text{ km s}^{-1}$  and a rotational velocity near zero. These results generally agree with our analysis when the group velocities of each cluster are averaged all together (Fig. 4). It is worth noting that the EVN observations have a beam size twice smaller than the JVN observations and four times better spectral resolution, which eventually allow us to obtain proper motion measurements about three times more accurate than the JVN measurements. On the one hand, since we measure a constant radial motion between radii of 300 au and 800 au, this finding explains why [Sugiyama et al.](#) inferred a similar infalling velocity at constant radius. On the other hand, by removing the constraint of constant radius, we are able to distinguish different tangential velocities associated with the five maser clusters.

## 5. Conclusions

We tested the existence of an equatorial plane of gas motion around Cepheus A HW2 by measuring the three-dimensional velocity vectors of the 6.7 GHz CH<sub>3</sub>OH maser cloudlets. Maser velocities are measured with an accuracy of a few  $0.1 \text{ km s}^{-1}$ . Our results can be summarized as follows.

1. We showed that the motions of the CH<sub>3</sub>OH maser cloudlets are dominated by a planar component. Its plane is inclined by  $26^\circ$  to the line of sight and intersects the plane of the sky along a line at a PA of  $134^\circ$ . The orientation of this plane coincides within about  $10^\circ$  with that of the core emission imaged by [Patel et al. \(2005\)](#), after assuming that dust and high-density gas move on circular orbits. This result supports the interpretation that the core emission surrounding HW2 represents a continuous structure and is not a projection of different hot cores along the line of sight.
2. The three-dimensional velocity field of the CH<sub>3</sub>OH masers shows that the CH<sub>3</sub>OH gas is contracting toward HW2 at a velocity of about  $2 \text{ km s}^{-1}$ , between radii of 300 and 900 au from HW2. This velocity component implies an accretion rate on the order of  $10^{-4} M_\odot \text{ yr}^{-1}$  at a distance of 300 au from the star. A planar infall toward HW2 in turn supports an accretion disk.
3. At a radius of 300 au from HW2, we measure a rotational component of about  $4 \text{ km s}^{-1}$ . The simple assumption of pure Keplerian rotation would result in a central mass of  $5\text{--}6 M_\odot$ , which would account for only a minor fraction of the bolometric luminosity of the HW2 object. If the rotational velocities measured with the masers were sub-Keplerian by a factor as small as 1/4, as can be expected under the influence of a weak magnetic field, that could reconcile the values of the (kinematic) stellar mass and luminosity. The presence of a magnetic field is supported by the recent detection of an ordered dust polarization field surrounding HW2, and can be tested with the 6.7 GHz maser observations once a new model for the polarized emission of this maser transition will be available.

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## Appendix A: Additional table

**Table A.1.** Parameters of the 6.7 GHz CH<sub>3</sub>OH maser cloudlets detected toward Cepheus A HW2.

Feature #	$V_{\text{LSR}}$ (km s <sup>-1</sup> )	$F_{\text{peak}}$ (Jy beam <sup>-1</sup> )	$\Delta x$ (mas)	$\Delta y$ (mas)	$V_x$ (km s <sup>-1</sup> )	$V_y$ (km s <sup>-1</sup> )
Cluster I: mean LSR velocity of $-2.31 \text{ km s}^{-1}$						
1	-1.68	1.04 (3 <sup>rd</sup> )	$19.38 \pm 0.17$	$-4.85 \pm 0.10$	...	...
2	-1.73	22.84 (1 <sup>st</sup> )	$-0.00 \pm 0.01$	$-0.01 \pm 0.01$	$0.61 \pm 0.06$	$1.49 \pm 0.06$
			...	...	...	...
			...	...	...	...
3	-1.90	0.57 (1 <sup>st</sup> )	$-25.36 \pm 0.18$	$4.35 \pm 0.09$	$1.53 \pm 0.37$	$0.64 \pm 0.18$
4	-2.03	1.81 (1 <sup>st</sup> )	$-51.33 \pm 0.06$	$17.08 \pm 0.05$	...	...
5	-2.17	4.07 (1 <sup>st</sup> )	$-47.75 \pm 0.06$	$23.57 \pm 0.05$	$-0.40 \pm 0.20$	$0.08 \pm 0.16$
6	-2.30	52.18 (1 <sup>st</sup> )	$-38.97 \pm 0.01$	$32.19 \pm 0.01$	$-0.10 \pm 0.07$	$1.40 \pm 0.07$
			...	...	...	...
			...	...	...	...
7	-2.30	3.09 (1 <sup>st</sup> )	$-63.16 \pm 0.15$	$47.39 \pm 0.12$	...	...
8	-2.47	32.57 (1 <sup>st</sup> )	$-76.76 \pm 0.02$	$54.28 \pm 0.02$	$-0.20 \pm 0.08$	$0.25 \pm 0.07$
			...	...	...	...
			...	...	...	...
9	-2.47	24.08 (1 <sup>st</sup> )	$-30.86 \pm 0.02$	$26.70 \pm 0.02$	$0.53 \pm 0.15$	$0.83 \pm 0.13$
			...	...	...	...
10	-2.65	6.43 (1 <sup>st</sup> )	$-92.35 \pm 0.06$	$62.36 \pm 0.04$	...	...
			...	...	...	...
11	-2.82	7.65 (1 <sup>st</sup> )	$-108.34 \pm 0.02$	$66.52 \pm 0.01$	$-0.19 \pm 0.08$	$0.24 \pm 0.07$
			...	...	...	...
			...	...	...	...
12	-2.91	0.94 (1 <sup>st</sup> )	$40.11 \pm 0.04$	$3.54 \pm 0.04$	$1.18 \pm 0.13$	$0.79 \pm 0.09$
13	-2.91	0.20 (3 <sup>rd</sup> )	$33.83 \pm 0.27$	$-2.56 \pm 0.17$	...	...
14	-3.00	0.25 (1 <sup>st</sup> )	$-121.31 \pm 0.10$	$67.99 \pm 0.10$	...	...
15	-3.00	0.15 (3 <sup>rd</sup> )	$19.37 \pm 0.18$	$-3.01 \pm 0.10$	...	...
Cluster II: mean LSR velocity of $-2.44 \text{ km s}^{-1}$						
16	-0.10	0.08 (3 <sup>rd</sup> )	$378.41 \pm 0.16$	$-135.89 \pm 0.13$	...	...
17	-0.45	0.78 (3 <sup>rd</sup> )	$311.09 \pm 0.02$	$-130.28 \pm 0.03$	...	...
18	-0.45	0.08 (3 <sup>rd</sup> )	$321.51 \pm 0.19$	$-140.62 \pm 0.26$	...	...
19	-0.94	0.13 (1 <sup>st</sup> )	$174.92 \pm 0.14$	$-107.37 \pm 0.13$	...	...
20	-1.20	0.21 (3 <sup>rd</sup> )	$323.52 \pm 0.07$	$-111.97 \pm 0.09$	...	...
21	-1.33	0.07 (3 <sup>rd</sup> )	$311.68 \pm 0.18$	$-107.07 \pm 0.17$	...	...
22	-1.81	1.67 (3 <sup>rd</sup> )	$368.78 \pm 0.08$	$-108.69 \pm 0.07$	...	...
23	-2.17	1.65 (3 <sup>rd</sup> )	$376.92 \pm 0.07$	$-112.75 \pm 0.08$	...	...
24	-2.42	9.42 (3 <sup>rd</sup> )	$359.55 \pm 0.09$	$-116.89 \pm 0.09$	...	...
25	-2.61	7.59 (3 <sup>rd</sup> )	$350.78 \pm 0.05$	$-117.32 \pm 0.04$	...	...
26	-2.69	1.67 (3 <sup>rd</sup> )	$360.80 \pm 0.13$	$-125.16 \pm 0.11$	...	...
27	-2.78	1.25 (1 <sup>st</sup> )	$371.31 \pm 0.10$	$-128.49 \pm 0.10$	$-2.89 \pm 0.15$	$3.93 \pm 0.16$
28	-2.82	0.86 (1 <sup>st</sup> )	$377.87 \pm 0.11$	$-135.48 \pm 0.10$	$-3.03 \pm 0.25$	$3.50 \pm 0.22$
29	-3.18	0.78 (1 <sup>st</sup> )	$415.83 \pm 0.03$	$-117.00 \pm 0.03$	$-2.41 \pm 0.11$	$1.54 \pm 0.11$
30	-3.22	0.68 (1 <sup>st</sup> )	$406.50 \pm 0.04$	$-135.77 \pm 0.04$	...	...
31	-3.44	0.59 (1 <sup>st</sup> )	$405.01 \pm 0.09$	$-140.44 \pm 0.10$	...	...
Cluster III: mean LSR velocity of $-3.63 \text{ km s}^{-1}$						
32	-3.35	0.41 (1 <sup>st</sup> )	$-259.51 \pm 0.14$	$160.15 \pm 0.20$	...	...
33	-3.44	2.37 (1 <sup>st</sup> )	$-242.84 \pm 0.03$	$145.19 \pm 0.03$	$-0.15 \pm 0.08$	$0.00 \pm 0.09$
34	-3.52	0.37 (3 <sup>rd</sup> )	$-286.04 \pm 0.22$	$185.77 \pm 0.25$	...	...
35	-3.53	2.52 (1 <sup>st</sup> )	$-253.31 \pm 0.05$	$147.95 \pm 0.06$	$-0.87 \pm 0.10$	$0.61 \pm 0.13$
36	-3.61	0.83 (3 <sup>rd</sup> )	$-273.56 \pm 0.21$	$123.62 \pm 0.19$	...	...
37	-3.62	1.49 (2 <sup>nd</sup> )	$-340.85 \pm 0.09$	$255.21 \pm 0.12$	...	...
38	-3.66	10.84 (1 <sup>st</sup> )	$-305.94 \pm 0.03$	$189.99 \pm 0.02$	...	...

**Notes.** Each maser emission center is labeled by decreasing LSR velocity in Col. 1, and grouped within the five clusters defined in Fig. 2. Columns 2 and 3 report the LSR velocity and brightness of the brightest spot of each cloudlet at the first epoch of detection (in brackets). Columns 4 and 5 give the relative centroid position of each cloudlet and their uncertainties in the east and north directions, respectively. The absolute position of the reference spot, belonging to cloudlet 2, is reported in Table 1. Columns 6 and 7 give the projected components of the cloudlet proper motion along the east and north directions, respectively. More details are provided in Sect. 3.



Table A.1. continued.

Feature #	$V_{\text{LSR}}$ (km s <sup>-1</sup> )	$F_{\text{peak}}$ (Jy beam <sup>-1</sup> )	$\Delta x$ (mas)	$\Delta y$ (mas)	$V_x$ (km s <sup>-1</sup> )	$V_y$ (km s <sup>-1</sup> )
39	-3.66	5.38 (1 <sup>st</sup> )	-351.66 ± 0.04	264.29 ± 0.06	...	...
40	-3.66	3.02 (1 <sup>st</sup> )	-297.34 ± 0.09	195.69 ± 0.12	-0.28 ± 0.23	0.87 ± 0.27
41	-3.66	2.60 (1 <sup>st</sup> )	-304.12 ± 0.11	201.49 ± 0.09	...	...
42	-3.70	1.45 (1 <sup>st</sup> )	-313.51 ± 0.20	196.58 ± 0.22	...	...
43	-4.49	0.10 (2 <sup>nd</sup> )	-183.38 ± 0.30	124.61 ± 0.19	...	...
Cluster IV: mean LSR velocity of -4.05 km s <sup>-1</sup>						
44	-3.88	5.88 (1 <sup>st</sup> )	1326.16 ± 0.04	-142.10 ± 0.03	...	...
45	-3.88	3.42 (1 <sup>st</sup> )	1329.83 ± 0.13	-153.46 ± 0.09	...	...
46	-3.92	5.28 (1 <sup>st</sup> )	1294.00 ± 0.06	-170.31 ± 0.06	...	...
47	-4.01	19.93 (1 <sup>st</sup> )	1312.58 ± 0.03	-143.87 ± 0.02	-0.44 ± 0.09	-0.03 ± 0.08
48	-4.01	5.28 (1 <sup>st</sup> )	1302.81 ± 0.09	-163.01 ± 0.07	...	...
49	-4.05	13.62 (1 <sup>st</sup> )	1314.18 ± 0.06	-157.25 ± 0.03	...	...
50	-4.19	17.16 (1 <sup>st</sup> )	1309.41 ± 0.02	-156.94 ± 0.02	...	...
			...	...	...	...
51	-4.27	1.36 (1 <sup>st</sup> )	1279.61 ± 0.15	-176.55 ± 0.10	...	...
52	-4.32	3.19 (1 <sup>st</sup> )	1287.24 ± 0.05	-163.20 ± 0.04	-0.65 ± 0.26	-0.03 ± 0.17
Cluster V: mean LSR velocity of -4.54 km s <sup>-1</sup>						
53	-4.36	1.55 (1 <sup>st</sup> )	111.09 ± 0.10	680.53 ± 0.18	0.38 ± 0.18	-1.31 ± 0.35
54	-4.36	1.10 (1 <sup>st</sup> )	118.36 ± 0.19	665.06 ± 0.32	...	...
55	-4.41	1.30 (1 <sup>st</sup> )	102.88 ± 0.10	694.45 ± 0.16	...	...
56	-4.41	0.24 (3 <sup>rd</sup> )	94.65 ± 0.11	703.51 ± 0.18	...	...
57	-4.54	0.26 (1 <sup>st</sup> )	131.94 ± 0.17	518.21 ± 0.17	1.07 ± 0.55	-3.35 ± 0.56
58	-4.58	0.90 (1 <sup>st</sup> )	122.06 ± 0.08	527.51 ± 0.08	0.59 ± 0.19	-1.47 ± 0.20
59	-4.67	0.51 (1 <sup>st</sup> )	102.48 ± 0.14	546.56 ± 0.12	0.27 ± 0.41	-2.27 ± 0.36
60	-4.71	1.78 (1 <sup>st</sup> )	92.14 ± 0.03	556.99 ± 0.04	1.32 ± 0.11	-2.03 ± 0.11
61	-4.80	0.29 (1 <sup>st</sup> )	86.64 ± 0.10	551.32 ± 0.09	1.75 ± 0.38	-0.96 ± 0.35
62	-4.80	0.29 (1 <sup>st</sup> )	73.17 ± 0.17	573.26 ± 0.16	1.57 ± 0.48	-2.21 ± 0.45
63	-4.93	0.18 (1 <sup>st</sup> )	-27.66 ± 0.16	551.60 ± 0.14	0.34 ± 0.59	-1.07 ± 0.43
64	-4.93	0.07 (1 <sup>st</sup> )	-35.71 ± 0.29	554.01 ± 0.37	...	...